Variation of H I Ly-α and O VI 103.2 nm in the Solar Corona from 1996 to 2000

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UVCS/SOHO measurements of H I Ly-α and O VI 103.2 nm intensities in the solar corona have been made from 1996 to the present, thus spanning the rising phase of cycle 23. At solar minimum, the corona consisted of large coronal holes at the poles and quiescent streamers at the equator. During the ascending phase of the cycle, the corona presented high latitude streamers. Finally, recent observations as the Sun approached solar maximum have shown the presence of coronal holes at the equator and streamers at the poles. Our observations provide descriptions of the various coronal structures that existed over the rising phase of the solar cycle. We compare spectral line intensities of quiescent equatorial streamers that occurred at solar minimum to those of high latitude and polar streamers observed toward solar maximum. We also compare spectral line intensities of solar minimum polar coronal holes to those of equatorial coronal holes present at solar maximum. We discuss how these results are related to the plasma properties.

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